continued observation. They have received it, and (with exception of those subsequent corrections in the numerical values which we have noticed and included in the above statement) they remain intact, and rank among the well-established facts of astronomy. Moreover, numerous other stars have been subjected to examination, some by one, some by the other method. And the result is not a little surprising. Up to the present time, out of all the stars examined, only a very few exhibit *any* distinctly measurable amount of parallax. The list hitherto accumulated consists only of about ten or at most a dozen. Of these α Centauri, in the southern hemisphere, is the nearest. It is a fine star of the first magnitude, the third or fourth in brightness of all the sidereal host. This is our next neighbour. On the other hand, Sirius, the brightest of all the stars, and Lyra (next to Sirius, one of the four most conspicuous stars in our hemisphere) stand low in the order of proximity. This, of course, only proves that among the stars there exists a very wide range of *absolute* brightness, but by no means invalidates the strong *à priori* reasons for admitting distance as a very important element in determining their relative *apparent* brightness.

(26.) But how, it will be asked, came such a seemingly insignificant object as this No. 61 to be selected for examination at all, to the exclusion or postponement of so many more conspicuous? We reply, by reason of its large apparent *proper motion*. None of the stars we see maintains *quite* the same relative situation among its compeers. It would be strange if it did. Unless nailed